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1978 Diffuse Auroral Boundaries and a Derived Auroral Boundary Index

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DMSP/F2 and F4 precipitating electron data are used to determine statistically the systematic variations of the equatorward boundary with Kp as a function of local time. The boundaries were chosen by hand for every DMSP/F2 satellite pass in 1978. These in turn are used to assess an algorithm developed to choose the boundaries automatically. From the statistical variations each boundary is projected to a midnight boundary. The projected midnight boundary served as an index of auroral activity-The Auroral Boundary Index. Listings of the 1978 hand- and computer-chosen

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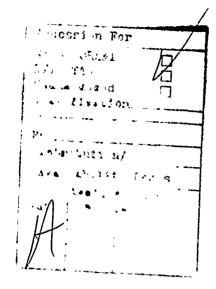
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Preface

We would like to thank Oracio Barbosa, Rebecca Carovillano, Deborah Gustafson, Robert Hilmer, Joan Hogan and Timothy Schumaker for their help in the production of the final hand-chosen bourdary set. Thanks also to Joseph Cronin and Dianne Riehl who developed the computer software which was used to produce the monthly plots, and to Kenneth McGee for preliminary plot production. A special second thanks to Dianne Riehl who, with consistent good humor, produced several "final" versions of all the monthly plots used in this report. Finally, another special thanks to Mary Outwater for secretarial support. The work of M.S. Gussenhoven and N. Heinemann was supported by the Air Force Geophysics Laboratory under Contract F19628-81-K-0032; the work of E. Holeman was supported by the Air Force Geophysics Laboratory Contract F19628-82-K-0039.

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1978 Diffuse Auroral Boundaries and a Derived Auroral Boundary Index

1. INTRODUCTION

This report has two aims. First, to provide the scientific community a compilation of all equatorward boundaries of the auroral oval as determined from the SSJ/3 data on the DMSP/F2 satellite for the year 1978. Second, to review and extend the knowledge of the systematics of the position of the equatorward boundary with respect to geomagnetic activity and the degree to which the boundary determination can be done by computer. From this we are able to construct a new index of auroral activity.

The first of these aims is motivated by research that has shown the boundary to move in a consistent and systematic manner in response to geomagnetic activity as measured by Kp, the velocity of the solar wind, and the strength of the north-south component of the Interplanetary Magnetic Field. ¹⁻⁷ This work has shown that the boundaries can be used as an indirect measure of the strength and orientation of the magnetospheric electric field when the assumption is made that the equatorward boundary maps to the zero energy Alfven layer in the magnetic equatorial plane. Calculations of the total cross magnetospheric potential drop determined from the average location of the boundaries were found to be in reasonable

⁽Received for publication 27 December 1982)

⁽Due to the large number of references cited above, they will not be listed here. See References, page 59.)

agreement with measurements of the average cross polar cap potential measured by probes. ⁷ Since the location of the boundary appears to reflect large-scale processes taking place in the magnetospheric system, it should provide a valuable tool in the study of geomagnetic phenomena.

The second of these aims is motivated by the need of the Global Weather Central of the Air Force Weather Service for a means of specifying in near real time the global extent of electron precipitation and the level of geomagnetic activity. The present work contributes to fulfilling the need of GWC in three ways. First, the set of boundaries compiled for this report represents a significant portion of the data set used to derive the systematics of the boundary location. Knowledge of these systematics has been used to extrapolate the global position of the boundary based on a single point determination of the boundary made in near real time from DMSP satellite data. 8 Second, the computer-chosen boundaries and the algorithm used in their selection represent an extension of work previously done for GWC on techniques for automatically determining the location of the boundary in near real time using the raw data from the SSJ/3 sensor on the DMSP satellites. 9 Lastly, the report shows how the boundary measurements made at different local times can be normalized to magnetic midnight to produce a new index of geomagnetic activity which correlates well with Kp. This new index will be directly available from the DMSP SSJ/4 data.

The report is divided into seven sections, including two appendices. Section 1 describes the instrumentation which provided the data used in this study. Section 2 deals with the format of the data used in the hand determination of the boundaries, the method used in the hand selection of the boundaries and the sources of error or ambiguity in these determinations. Section 3 describes the computer algorithm used to choose the boundaries. Section 4 provides an analysis of the discrepancies between the hand and computer techniques. Section 5 describes the process by which the auroral activity index was produced. Appendix A gives plots of the equivalent midnight boundary from the hand determined data set by month for 1978; the Auroral Boundary Index for one year. Appendix B gives a complete listing of both boundary data sets for 1978.

^{8.} Hardy, D.A., and Holeman, E. (1983) The Global Auroral Boundary Code for the Global Weather Central of the Air Weather Service (to be published).

^{9.} Hardy, D.A., and MacKean, R. (1980) An Algorithm for Determining the Boundary of Auroral Precipitation Using Data from the SSJ/3 Sensor, AlfGL-TR-80-0028, AD A084482.

2. INSTRUMENTATION

DMSP/F2, a three-axis stabilized satellite, was launched into a near sunsynchronous, circular orbit at an altitude of 840 km in June 1977. Its orbital period was 101 min; the nominal inclination, 98.75°. At launch the orbit was centered near the 0700-1900 meridian but was subject to a very slow precession toward later local times. Due to the offset between the Earth spin axis and magnetic axis, the orbit had significant diurnal and seasonal variations in the magnetic local timemagnetic latitude frame of reference. Thus, equatorial auroral boundaries could be determined over a wider range of MLT than might be assumed from the restricted geographical local time locations of the orbit. Figure 1 shows the diurnal coverage

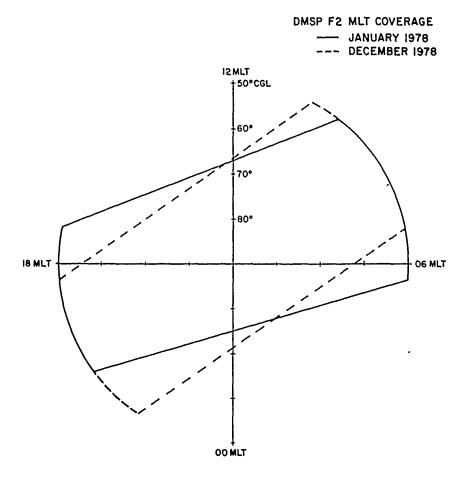


Figure 1. A Polar Plot in Magnetic Local Time-Magnetic Latitude Coordinates Showing the Diurnal Range of the DMSP/F2 Orbits in January 1978 (solid line), and in December 1978 (dashed line)

of DMSP/F2 in magnetic local time and magnetic latitude for January 1978 (solid line), and December 1978 (dashed line). Good coverage exists in the dawn and dusk sectors. Little coverage exists in the noon, post-noon and the midnight, post-midnight sectors. To extend the boundary statistics into these regions data from a later satellite, DMSP/F4, launched in April 1979 was also used. Identical particle detectors were flown on both satellites. It should also be noted that DMSP satellites are operational Air Force satellites. As such, except during periods of down-link transmissions, data are almost always being recorded. In 1980-1982, because of the failure of DMSP/F5 at launch, there is a break in the coverage, but DMSP/F6 was launched in late 1982.

The particle detector on DMSP/F2 consists of two curved plate electrostatic analyzers that measure the fluxes of electrons in 16 energy channels between 50 eV and 20 keV once per second. The apertures of the analyzers always face in the local zenith direction such that at auroral and polar cap latitudes they detect precipitating rather than backscattered and/or trapped electrons. One analyzer covers the energy range from 50 eV to 1 keV with a geometric factor of 4×10^{-4} cm² ster and a $\Delta E/E$ of 13 percent. The other analyzer covers the energy range from 1 keV to 20 keV with a geometric factor of 10^{-3} cm²-ster and a $\Delta E/E$ of 9 percent. The large geometric factors insure that the flux level for electrons in the diffuse aurora is well above the detector's sensitivity. A detailed description of the detector is given by Hardy et al. 10

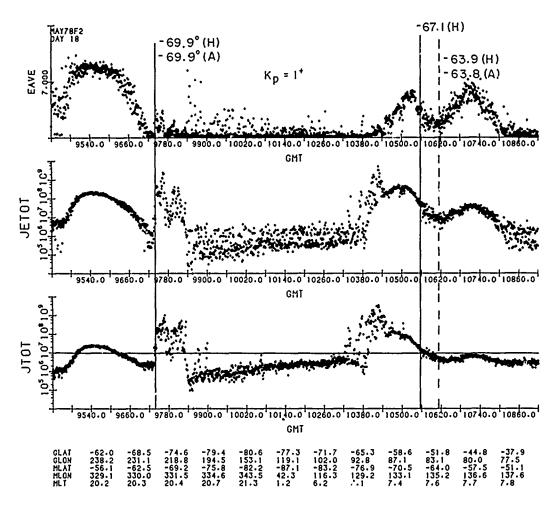
3. HAND SELECTION OF EQUATORWARD BOUNDARIES

An example of DMSP/F2 electron data, taken from a south polar pass on 18 May 1978, is given in Figure 2. For this pass Kp = 1+. Data are plotted as JTOT, the directional integral flux (cm²-sec-ster)⁻¹ in the bottom panel; JETOT, the directional energy flux (keV/cm²-sec-ster) in the middle panel; and EAVE, the average energy in keV in the top panel. The scale for EAVE is linear. These quantities are plotted as functions of universal time in seconds of the day, the geographic and corrected geomagnetic latitudes and longitudes, and the magnetic local time of the satellite all projected to an altitude of 110 km.

Several features of the electron precipitation that pertain to the choice of auroral boundaries are illustrated in Figure 2, First, equatorward of the auroral precipitation there is a broad region over which there is a slight rise in JTOT with relatively large values of JETOT and EAVE. In Figure 2 these lie between

^{10.} Hardy, D.A., Gussenhoven, MS., and Huber, A. (1979) The Precipitating Electron Detectors (SSJ/3) for the Block 5D/Flights 2-5 DMSP Satellites: Calibration and Data Presentation, AFGL-TR-79-0210, AD A083136.

9480-9690 UT and between 10,620-10,800 UT. The increases are due to radiation belt particles that penetrate the detector casing and directly stimulate the channel-trons. Since they have nothing to do with the auroral precipitation, they must be differentiated from the auroral electrons when determining boundaries. Due to the difference in the size of channeltrons used in the two detectors, radiation belt contamination is largely limited to the energy channels of the 1 to 20 keV detector.



Figur. 2. Integral Flux in (cm²-ste·-s)⁻¹ (bottom panel), Energy Flux in keV (cm²-ster-s)⁻¹ (middle panel), and Average Energy in keV (top panel) of Precipitating Electrons Measured by the DMSP/F2 Satellite Passing Over the South Pole on 18 May 1978. These values are plotted as functions of universal time (in seconds), geographic and corrected geomagnetic latitudes and longitudes and the magnetic local time of the satellite all projected to an altitude of 100 km. Solid (dashed) vertical lines indicate 10⁷ ("better") equatorward boundaries of precipitating auroral electrons chosen by hand (H). Algorithm-chosen boundaries (A) are indicated by a dot-dash line when they differ from the hand boundary by more than 0.3° CGL

Second, we have found the condition $JTOT > 10^7 \text{ (cm}^2\text{-sec-ster)}^{-1}$, indicated in Figure 2 by the horizontal line (the 10^7 level) provides a useful "zero-order" criterion for selecting the equatorward boundaries of the oval. The solid, vertical lines drawn where the data points rise above the 10^7 level, going in the poleward direction, are referred to as 10^7 boundaries.

Third, the evening equatorward boundary is sharper than its morning side counterpart. On the morning side in Figure 2 JTOT rose from background to $10^7 \, (\mathrm{cm}^2\text{-sec-ster})^{-1}$ over 3.2° latitude, whereas the evening side rise was nearly instantaneous. Generally, gradients in the electron flux with latitude are found near both equatorward boundaries. The gradient never extends more than a few degrees latitude in the evening sector, but can cover as much as 10° on the morning side where it is always perceptibly present. A dashed line on the morning side in Figure 2 is drawn at the point where JTOT rises noticeably above background, to include the flux gradient within the auroral boundary. This is referred to as a "better" boundary.

Because of the nearly instantaneous rise of the data on the evening side, most of the evening boundaries are 10⁷ boundaries. Because of flux gradients on the morning side, morning boundaries are generally "better" boundaries. In Appendix B "better" boundaries (or 10⁷ boundaries when the two are the same) are listed. This is the set of boundaries which has been used for various statistical studies.

The 10⁷ boundary is quite precise, but the criteria used for choosing the "better" boundary are more subjective, and cause some ambiguity. On both morning and evening sides of the oval, the onset of electron precipitation can be obscured in various ways.

In the evening sector the principal ambiguity in choosing boundaries is caused by contamination of the low energy channels of the detector by photoelectrons. Gussenhoven et al 6 discuss this problem in detail. In these data this contamination shows up as a high equatorward background level which can exceed a level of 0.5×10^7 on the JTOT scale. Although the choice of auroral boundaries is generally clear the ambiguity results from the fact that it is impossible to .ell if, and how far, the aurora might extend below the JTOT level produced by the photoelectrons. Clearly, when the background level is close to or above the 10^7 level on the JTOT scale choosing a 10^7 boundary is impossible.

In the morning sector uncertainties in identifying the position of the boundary are more severe. There are three sources of uncertainty: (1) a much more gradual latitudinal onset of precipitation than is found in the evening, (2) overlap between regions of energetic electron precipitation and radiation belt contamination of corresponding energy channels, and (3) the existence of energetic plasma close to but detached from the boundary.

The gradual latitudinal onsets of precipitation with increasing latitude, are called ramps (Figure 2, morningside). Ramps may occur on both morning and evening sides of the north and south poles, but are less frequent and less extensive on the evening side. They occur nearly all the time on the morning side generally producing a lower boundary when compared to the evening side. Ramps cause severe problems in two different ways: (1) when they are combined with a high equatorward background (as discussed in the case of evening boundary ambiguities), and (2) when the morning side gradient extends into regions of radiation belt contamination. Even when not obscured there can be ambiguity in choosing where a gradual ramp begins. In choosing the morning boundaries it is almost always necessary to use changes in the average energy and the energy flux, as well as those in the number flux.

Overlap between the region of energetic electron precipitation in the auroral zone and the region of contamination of energy channels by radiation belt particles is shown in Figure 3, a pass over the south pole on 4 April 1978, for which Kp = 5-. On the morning side of the oval one can see that the rise in JTOT, JETOI and EAVE, due to contamination, is contiguous to the region of auroral precipitation. When the radiation belt particles exhibit a characteristically smooth curve an attempt is made to determine the onset of auroral particle precipitation by drawing the boundary at the point where irregularity begins as was done in this example. In this case the morning boundary was drawn at -60.7° MLAT, but choices down to -56.6° (the 10⁷ boundary) are defensible giving an uncertainty of 4.1°. However, when the radiation belt and particle signatures are both irregular there is even greater uncertainty in the boundary choice. Because of the sudden onset of auroral electron precipitation on the evening side, radiation belt particle contamination is rarely a problem, as Figure 3 also shows.

At lower Kp, as in Figure 2, radiation belt interference produces no problems since the aurora is contracted enough to leave the radiation belt particles clearly outside the oval. Radiation belt interference is most problematic at moderately high activity levels (Kp = 3 to 6). For these activities it is often the case that a clear separation between JTOT levels produced by auroral and radiation belt effects cannot be made. At very high Kp, the problems are again manageable since fluxes are so intense they appear to mask the effect of the radiation belt particles. An example of this type of occurrence is snown in Figure 4, a north pole pass on 28 August 1978, for which Kp = 8.

The third effect (detached energetic plasma) can be found in both morning and evening regions and an example is shown in Figure 5. This example shows a north pole pass on 21 January 1978, during a time of magnetospheric quiet (Kp = 0). On the evening side there are small regions of energetic electrons near to but detached

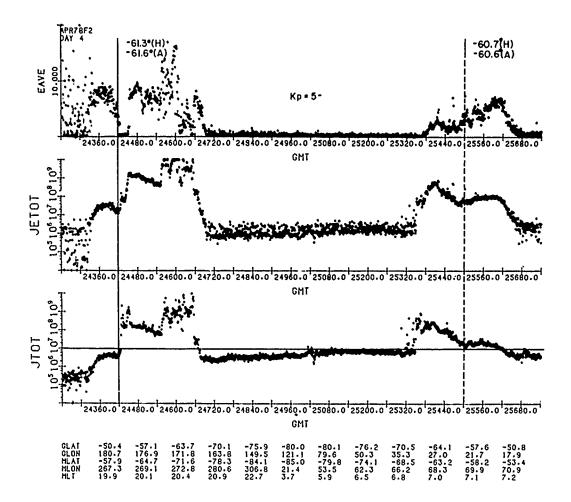


Figure 3. Pracipitating Electron Data for a South Pole Pass on 4 April 1978. The format is the same as that in Figure 2

from the more poleward precipitation region. These can occur during active times, as well, and are ignored in selecting the boundaries. Ambiguity may result when these regions are very broad and/or very close to the auroral region. (Note that Figure 5 is also an example of high equatorward background from photoelectrons from the sunlit conjugate hemisphere, as discussed above.)

The final problem encountered in this data set is with those morning south pole equatorward boundaries which fall between 0900-1300 MLT. Electron precipitation patterns in the diffuse auroral region during these local times are observed to vary from a normal diffuse aurora to the complete disappearance of the diffuse aurora within detector sensitivity. In the latter instances only the magnetospheric cusp is encountered. An example of a pass with no dayside diffuse aurora is shown in

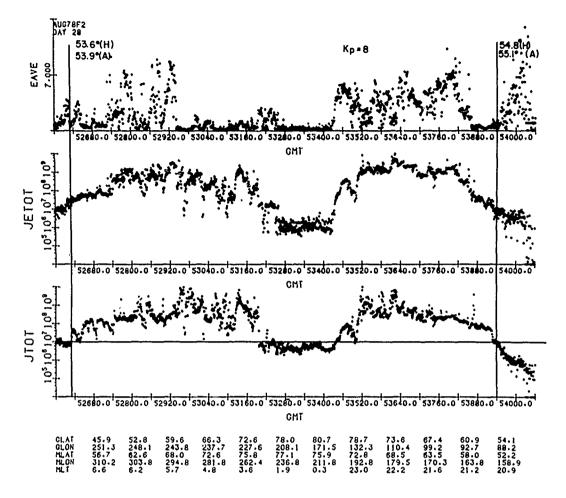


Figure 4. Precipitating Electron Data for a North Pole Pass on 28 August 1978. The format is the same as that in Figure 2

Figure 6, a south pole pass on 12 December 1978, for which Kp = 2-. For cases such as these no equatorward boundary is listed. Problems arise when the diffuse aurora is weak and irregular. As may be seen in Figure 1, DMSP/F2 is only in the pre-noon sector for a few passes per day, so any problems with the 1978 data sets should be minimal. The dynamics of the dayside diffuse auroral electron precipitation will be discussed in a forthcoming paper. For the time being, however, listed boundaries between 0900-1300 MLT should be treated judiciously, and we advise referencing the actual data to determine what is being measured in these instances.

The above effects produced greater than 1 degree uncertainty in the equatorial boundary only for a small percentage of the total cases. On the evening side

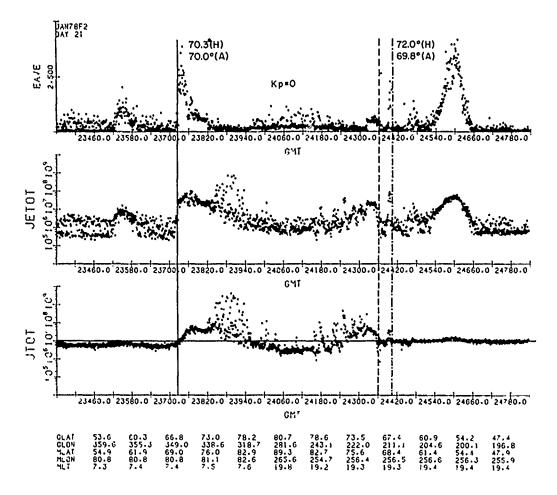


Figure 5. Precipitating Electron Data for a North Pole Pass on 21 January 1978. The format is the same as that in Figure 2

approximately 90 percent of boundaries were determined with no ambiguity; on the morning side, approximately 70 percent. After making more than 20,000 boundary determinations by hand, we feel that a high enough level of consistency has been reached to justify using the hand-chosen boundary set to verify the algorithm which has been developed to select boundaries automatically. In actuality, this verification turned out to be a two way process which has enabled us to decide which of several algorithm tests works best for determining morning and evening auroral boundaries (as discussed in Section 4), and to eliminate from the hand-chosen data set most of the glaring errors. As a result of this cross-verification, we have two data sets in which we have a high level of confidence.

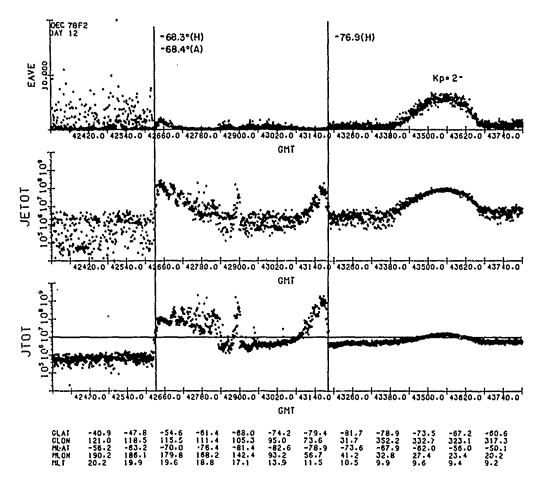


Figure 6. Precipitating Electron Data for a South Pole Pass on 12 December 1978. The format is the same as that in Figure 2

4. COMPUTER ALGORITHM FOR BOUNDARY SELECTION

In devising methods for choosing equatorward auroral boundaries by computer we made use of two factors that we found to be generally applicable in choosing the boundaries by hand; (1) on the evening side the electron precipitation begins (on the equatorward edge proceed ng poleward) with increases in the flux at the lowest energies, often resulting in a gradual rise in average energy, and (2) on the morning side the precipitation onset is most reliably seen in the 1-10 keV energy range, resulting in an average energy rise which is rather abrupt. Thus, an increase over background in the energy range < 1 keV will be used to signal the evening boundary, and an increase over background in the keV electrons, used to signal the morning

boundary. The complexity of the tests used reflects the difficulty often encountered in determining that the signal is above background.

Photoelectron contamination is avoided in the low energy particle signal by eliminating counts from the two lowest energy channels. Contamination of the higher energy signal by radiation belt particles is more complex and requires tests that make use of the fact that the low energy detector (channels 9 to 16) has only ~1/10 the contamination of the high energy detectors (channels 1 to 8), and that the contamination in the high energy detector gives a uniform count signal in all channels.

In this section we will describe the preparation of the data base, the construction of test quantities from the data, and the application of the tests that produced the 1978 set of computer- (algorithm-) chosen equatorward boundaries. The DMSP test data for the SSJ/3 instrument are stored on a series of 6250 BPI tapes, each containing several months of raw data from either the F2 or F4 satellite. The data base is analyzed one month at a time for boundary information. The first step in this procedure is to run program DSFMEP on a month's data with control parameters set to strip off all data with absolute magnetic latitude greater than 50°. Program STDKC3 then searches this intermediate file for boundary crossings. STDKC3 logically divides the file into quarter orbit segments by successive calls to subroutine SEARCE. For each call to SEARCE, information for up to 300 consecutive 4-sec data records is stored in the matrix ITST for subsequent analysis. For each quarter orbit, SEARCE determines whether the data are occurring in ascending or descending order of absolute geomagnetic latitude as a function of time. If the geomagnetic latitude is decreasing, the data are stored in inverse order beginning in location 300 so that STDKC3 always sees a quarter orbit file in order of increasing geomagnetic latitude.

The ITST matrix is preset to a value of -1. and the relative location in ITST for a particular 4-sec data record is calculated from the elapsed universal time since the beginning of the quarter orbit. Hence, time gaps are automatically embedded as a series of one or more lines of all minus ones. The end of a quarter orbit is defined to be where the data goes from one side of geomagnetic midnight (43, 200 secs) to the other for the ascending latitude half of a polar pass and the 50° latitude boundary for the descending magnetic latitude half. The last valid position in the data file is marked for the next call and control returned to STDKC3.

The DMSP raw data file is a series of consecutive data records, each containing a time and ephemeris record and four consecutive 16-channel spectral observations seen by the SSJ/3 instrument. Each spectrum is acquired over approximately one second of elapsed universal time. For statistical purposes these four spectra are added together to form a unit data record and its information stored in one line of ITST. ITST is dimensioned as (300, 18) and the 18 words corresponding to the N'th ITST line are defined as follows:

ITST (N, 1) through ITST (N, f of counts in consecutive energy channel pairs 1+2, 3+4, ..., 15+16 (to el numbers increase with decreasing energy;

ITST (N, 9): Counts in channel 8 (high energy detector, 985 eV channel);

ITST (N, 10): Counts in channel 9 (low energy detector, 972 eV channel);

ITST (N, 11): Universal time;

ITST (N, 12): Geomagnetic latitude;

ITST (N, 13): Magnetic local time;

ITST (N, 14): 110 km latitude;

ITS1 (N, 15): 110 km longitude;

ITSI (N, 16): Altitude;

ITST (N, 17): Geographic latitude;

ITST (N, 18): Geographic longitude.

STDKC3 analyzes the 300 record ITST quarter orbit file in sequential order until three of six defined pseudo statistical tests are satisfied. When this happens a successful boundary determination has been made. The universal time at which each test was satisfied along with its corresponding magnetic latitude and local time are recorded in the computer boundary file. The six test quantities are defined as follows:

JTST1: low energy integral count signal omitting the two lowest energy channels.

$$JTST_1 = S_1 \sum_{i=9}^{14} N_i$$
,

where N; is the observed counts in channel i, and

 $S_1 = 0.1$ (scale factor).

JTST₂: 8-4-2-1 weighted (toward high energy) RMS variance of 1st four channel pairs.

$$\operatorname{JTST}_2 = \sqrt{\frac{s_2}{T}} \quad \left\{ \frac{1}{w} \quad \sqrt{w \quad \sum_{i=1}^4 w_i I_i^2 - \left(\sum_{i=1}^4 w_i I_i\right)^2} \right\} \ ,$$

where

$$w = \sum_{i=1}^{4} w_i;$$

 $w_i = \{8, 4, 2, 1\}$ weighting factors;

 $I_i : \cdots$ ats in i'th channel pair, 1+2, 3+4, and so on;

$$\overline{I} = \frac{\sum_{i=1}^{4} w_i I_i}{w};$$

$$S_2 = 10.$$

 ${
m JTST}_3$: 1-2-4-8 weighted (toward low energy) rms variance in first four channel pairs. The definition is the same as for ${
m JTST}_2$, but with

$$w_i = \{1, 2, 4, 8\}$$
;

$$S_3 = 10$$
.

 ${
m JTST_4}$: difference function between 5.5-20 keV counts and 1-3.5 keV counts.

$$JTST_4 = \frac{S_4 \mid I_1 + I_2 - I_3 - I_4 \mid}{\sqrt{\max(I_1 + I_2, I_3 + I_4)}}$$

where

$$S_4 = 10$$
.

 ${
m JTST}_{\xi}$: unweighted rms variance in 1st four channel pairs. The definition is the same as for ${
m JTST}_2$, but with

$$w_i = \{1, 1, 1, 1\}$$
;

$$S_5 = 10$$
.

JTST6: ratio of counts in channel 8 to those in channel 9.

$$JTST_6 = S_6 \frac{N_8}{N_9};$$

$$S_6 = -10$$
.

Tests 2, 3 and 5 are of the form $S\sqrt{1}\sigma_p$ where σ_p is the normalized statistical variance $(\sigma_p = \frac{\sigma}{1})$. It was found empirically that this form gave a more reliable signal than either σ or σ_p alone.

JTST, is the sum of the counts in the six channels of the SSJ/3 sensor between 110 eV and 1000 eV. On the evening side of the oval counts in this energy range are normally observed first as the satellite passes the auroral oval boundary. The two lowest channels are excluded since they are often contaminated by conjugate photoelectrons. JTST2 is the weighted rms variance of the four channel pairs in the energy range from 1 to 20 keV. These are the channels that are affected by penetrating particles from the radiation telts. Since these particles directly stimulate the channeltrons each of the eight channels should see the same number of counts within statistics such that the rms variance (JTST₅) is small when there is no other signal. If counts from auroral fluxes are present they should be highly non-uniform over this energy range and would act to increase the rms variance. The weighting for JTST, is toward the highest energy channels. JTST, is the same as JTST₂ except that the weighting is toward the 1 keV end of the range. $JTST_A$ is the absolute difference in the 4-sec sum of counts in the four channels between 5.5 and 20 keV and the four channels between 1 keV and 3.5 keV divided by the square root of whichever of these sums is greater. The idea again is that if auroral counts are present this difference will be larger than if only counts from penetrating particles are present. This test was not used to determine the final boundary but was retained since it is the same as the test developed by Hardy and MacKean for GWC. ${
m JTST}_5$ is the same as ${
m JTST}_2$ and ${
m JTSI}_3$ except with no weighting. JTST6 is the ratio of channels 8 to channel 9. Both these channels are at approximately 1 keV but channel 8 should see the effects of penetrating particles and channel 9 should not. Based on the difference in the geometric factors between channel 8 and 9 the ratio should be approximately 3 if both channels were measuring the same flux. If the ratio is much higher than this it means that channel 8 has significant counts due to penetrating electrons.

STDKC3 treats each of these six quantities as an independent statistic with an empirically determined normal, average, or threshold value when the satellite is at latitudes below the auroral boundary. When a statistic is above its threshold value for a sufficiently long time period, the boundary flag for that statistic is turned on. That time period was chosen as one 4-sec interval for JTST and three 4-sec intervals for the remaining tests. STDKC3 scans the quarter orbit file from its beginning, setting the flag for each test until flags are on for JTST, JTST, and JTST at the same time. STDKC3 retains the geomagnetic latitude and magnetic local time at which the flag for each test was set. Once the flags for JTST, JTST, and JTST, are set the boundary is chosen on the morning side of the oval at the point where the flag for JTST, was set and on the evening side of the oval at the point where JTST, was set.

Once a signal for a given test is turned on it is possible for its corresponding statistic to return to its threshold level or below. IF ${\rm JTST}_1$ and ${\rm JTST}_5$ along with with either ${\rm JTST}_2$ or ${\rm JTST}_3$ are below threshold for two consecutive 4-sec time intervals before the boundary search reaches a normal termination (${\rm JTST}_2$, 5, 6 on), all signals are turned off and the search resumed. Independent of that ${\rm JTST}_6$ is turned off if its statistic is below threshold for five consecutive 4-sec time intervals.

The actual thresholds chosen for the six quantities were $T_c = \{12, 18, 16, 30, 18, -50\}$. These were chosen empirically using approximately 500 case histories before the procedure was applied to the entire DMSP data base to construct the final computer boundary file.

Time gaps in the data file had to be dealt with. One and two 4-sec time gaps were ignored since they could cause at most a 8-sec error in the boundary determination. However, if a longer time gap occurred, all flags were turned off. Also, if any of the six signals occurred immediately after a time gap longer than 8 sec, its boundary was flagged as questionable in the computer boundary file. The beginning of a quarter orbit file was logically treated as if it were occurring after a large time gap for the time gap logic. If the data were exhausted before all flags were set, the current status of each was written into the boundary file, and labelled as questionable. Also, the boundary latitudes were examined and labelled questionable if they were above 72°. A record was written for each of the quarter orbit files, and analyzed for completeness with all questionable data flagged.

5. ANALYSIS OF DIFFERENCES BETWEEN HAND-CHOSEN AND ALGORITHM-CHOSEN BOUNDARIES

Two methods of analyzing the hand- and algorithm-chosen sets of boundaries are given in this section. The first method relates each set separately to geomagnetic activity, as measured by Kp, by finding, for each local time sector, a linear relationship between Kp and the boundaries. We then compare the linear relationships of both sets, and the residual scatter found within each set. The second method looks directly at the differences between the two boundary choices, and analyzes the causes of the differences when they are found to be systematic and large.

Relationship of boundary sets to Kp. To relate the boundary sets to Kp, 24 1-hr bins in magnetic local time were created and individual boundaries assigned to a bin based on their MLT. Each boundary was tagged with the value of Kp at the time of the boundary crossing. Within each local time bin a linear regression was performed on the corrected geomagnetic latitudes (λ) of the boundaries vs Kp of the form:

 $\lambda = \alpha + \beta \text{ Kp}. \tag{1}$

Table 1 is a list of the results of the regression for those MLT bins in which there are sufficient data. In the table we give the range of the MLT bin, the number of boundaries in the bin used in the regression, the intercept (α) and the slope (β) of the regression line and the correlation coefficient (cc). The results from the hand-chosen and computer-chosen boundaries are listed separately. Hand-chosen boundaries using both F2 and F4 satellite data between September 1977 and February 1980 were used to obtain the regression equations. Computer-chosen boundaries used F2 satellite data from 1978, with selected values from 1979 and 1980 to complete the midnight sector statistics.

Comparing the regression equations from the hand- and computer-chosen boundaries one notes several systematic differences. On the morning side of the oval (MLT's between 0400 and 1200) the regressions on the hand boundaries give intercepts roughly 1° higher and slopes approximated 0.4°/Kp more negative than the computer-chosen boundaries. On the evening side of the oval (MLT's between 1600 and 2400) the regressions on the hand boundaries give intercepts roughly 0.4° lower and slopes approximately the same as the computer-chosen boundaries results.

These differences reflect the limitations on the computer techniques. We attribute the differences on the morning side to the general problem the computer algorithm has in picking up the onset of morning side ramps and to the fact that this difficulty increases with increasing activity as the ramp region and the region of contamination by penetrating radiation more and more overlap. This problem results in computer boundaries higher than hand boundaries and a difference between the two that increases with activity. This leads to a higher average location for the computer boundaries at high Kp such that the slope of the regression curve is forced down and the intercept lowered. The evening side differences result from the threshold for the computer test being set at a high enough level to ensure that counts due to photoelectrons in the low energy channel would not trigger the test. This conservative approach results in the vast majority of computer boundaries being chosen at higher latitudes than the hand boundaries. Since the onset of precipitation on the evening side is normally abrupt at all levels of activity, on average, the difference between hand- and computer-chosen boundaries is the same resulting in the observed 0.4° offset but the same slopes.

Table 1. Regression Values for $\lambda = \alpha + \beta K_p$

	Hand-chosen	shosen				Compute	Computer-chosen		
	Number	α	В	၁၁	Number	σ	β	၁၁	
04-05	267	67.7	-1.48	-0.57	112	67.4	-1.28	-0.52	
90-90	1123	8.79	-1.87	-0.71	865	67.0	-1.57	-0.69	
06-07	2462	68.2	-1.90	-0.74	1670	67.2	-1.51	-0.70	
07 -08	3159	68.9	-1.91	-0.76	1926	68.0	-1.52	-0.72	
60-80	2159	69,3	-1.87	-0.73	1045	68.2	-1,45	-0.68	
09-10	1178	69, 5	-1.69	-0.66	441	68.5	-1.21	-0.62	
10-11	864	69, 5	-1.41	-0.57	220	69.0	-1.12	-0.60	
11-12	513	70.1	-1.25	-0.52	27	67.3	-0.47	-0.37	
16-17	204	71.6	-1.28	-0.66	123	72.0	-1.15	-0.63	
17-18	526	71.1	-1.31	-0.69	341	72.0	-1.35	-0.75	
18-19	997	71.2	-1.74	-0.82	652	71.6	-1.71	-0.80	
19-20	2469	70.4	-1.83	-0.82	1537	70.7	-1.85	-0.85	
20-21	3309	69.4	-1.89	-0.82	1934	69.8	-1.82	-0.83	
21-22	3092	68.6	-1.86	-0.73	1513	69.1	-1.73	-0.81	
22-23	1485	61.9	-1.78	-0.77	861	68.4	-1.52	-0.71	
23-24	461	67.8	-2.07	-0.81	179	67.7	-1.64	-0.54	

We next investigated the degree of scatter in the two data sets from their respective regression equations. This was done by subtracting the appropriate regression equation value from each boundary measurement and determining the frequency distribution of the residuals. In Table 2 are listed the percentages of this residual scatter within \pm 1°, \pm 2°, \pm 3° of zero for computer- and hand-chosen boundaries for morning and evening sides of the oval, both separately and together.

Table 2. Percentages of Residual Scatter

	Morning	Side	Evening	Side	Entire C	val
	Computer	Hand	Computer	Hand	Computer	Hand
± 1°	36%	30%	44%	40%	40%	36%
± 2°	64%	56%	75%	70%	68%	65%
± 3°	82%	76%	90%	86%	86%	82%

The spread in the unaccounted scatter for the computer- and hand-chosen boundaries is comparable, the computer boundary spread being approximately 5% narrower. This reflects a higher consistency in the computer-chosen over the hand-chosen boundaries. It is also the case that for both computer- and hand-chosen boundaries the spread is approximately 10° greater on the morning than on the evening side of the oval. Figures 7 and 8 are plots of the distribution of the residual scatter taken over the entire oval for the hand and computer boundaries, respectively. A bin size of 0.1° is used and the number normalized to 1000 cases. It is clear that the distributions are Gaussian. Based upon this and the results listed in Table 2 one can assign to the distributions a sigma of approximately 2° for both computer- and hand-chosen boundaries.

Direct comparison of hand- and computer-chosen boundaries. Distribution of differences between hand- and computer-chosen boundaries for the evening and morning sides of the oval, both separately and together, are shown in Figures 9 through 11. In addition, in Table 3 we show the percentage of the differences within \pm 1°, \pm 2° and \pm 3° of zero for the three cases and a comparison to the earlier work of Hardy and MacKean. 9

The figures and tables illustrate three points. First, all three distribution are strongly skewed towards negative values for the differences between computer-and hand-chosen boundaries. As discussed above, this arises due to the requirement of the computer algorithm for a clearer signal than that required by eye in order to consistently avoid the effects of background and noise in the selection

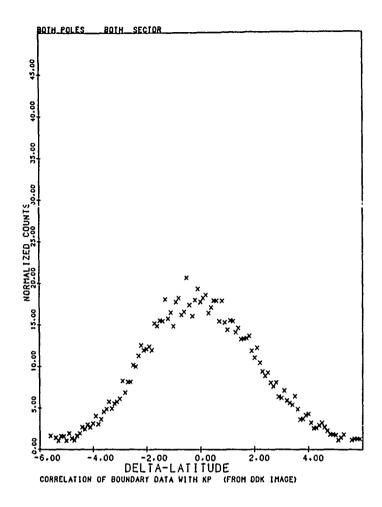


Figure 7. Distribution of the Residual Scatter in the Handchosen Boundary Set When Compared to Linear Regression Values Obtained From the Same Set. The distribution is accumulated in 0.1° bins and normalized to 1000 cases

process. Second, the distribution on the evening side of the oval is much narrower than on the morning side reflecting the more abrupt and therefore clearer onset of precipitation on the evening side. Lastly, the results are in very close agreement to those obtained by Hardy and MacKean⁹ which were derived using slightly different statistical tests.

We next analyze the source of differences between hand and computer boundaries that are greater than \pm 3° on the morning side and greater than \pm 2° on the evening side of the oval. The reasons for the larger differences between the hand-chosen and algorithm-chosen boundaries are given in Table 4. Morning and evening

sector discrepancies are listed separately because the number of cases differs. A discussion of these large differences in order of frequency of occurrence is given below.

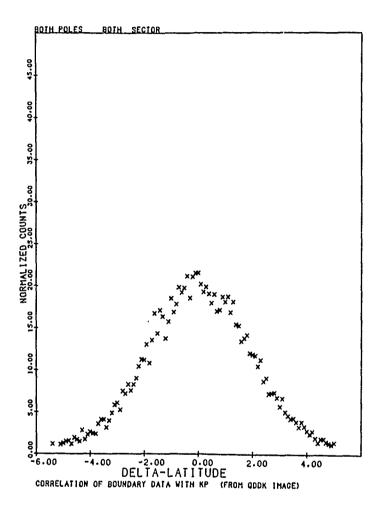


Figure 8. Distribution of the Residual Scatter in the Computer-chosen Boundary Set When Compared to Linear Regression Values Obtained From the Same Set. The distribution is accumulated in 0.1° bins and normalized to 1000 cases

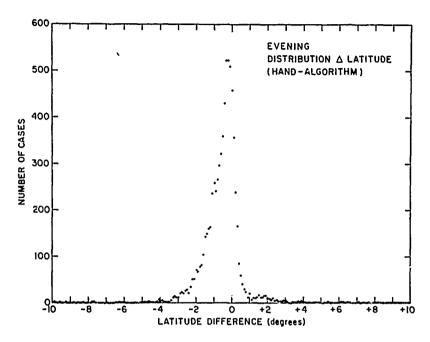


Figure 9. Distribution of the Differences Between Hand- and Computer-Chosen Boundaries $(\lambda_{II}$ - $\lambda_A)$ for the Evening Sector. The distribution is accumulated in 0.1° bins

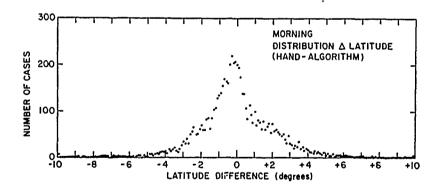


Figure 10. Same as Figure 9, for Morning Sector

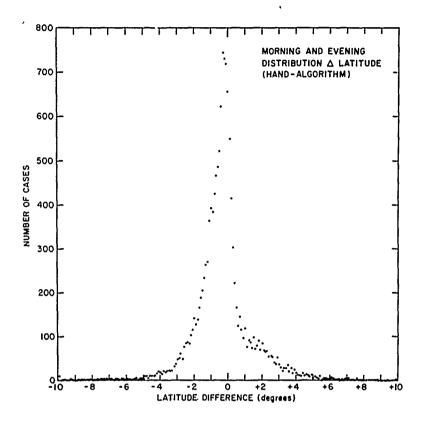


Figure 11. Same as Figure 9, for Evening and Morning Sectors Combined ${\bf S}$

Table 3. Percentages of Differences

	Evening	Morning	Composite	H-M Evening	H-M Morning
± 1°	74%	50%	63%	80%	50%
± 2°	93%	74%	84%	95%	67%
± 3°	98%	89%	94%	99%	82%

Table 4. Hand- vs Algorithm-Chosen Boundaries

	Morning				Evening		
Morning Discrepancies (> ± 3° CGM)				Evening Discrepancies (> ± 2° CGM)	N)		
Reason	Number	% of Total Morning Boundaries	% of Morning Discrepancies	Reason	Number	% of Total Evening Boundaries	% of Evening Discrepancies
Radiation Belt Interference	241	3.8	(34.7)	Evening Ramp	303	4.3	(64.3)
Irregularities Equatorward Edge	241	3.8	(34.7)	Irregularities Equatorward Edge	99	0.9	(14.0)
Spurs	7.1	1.1	(10.2)	Spurs	36	0.5	(7.6)
Magnetic Local Noon Criteria	54	0.8	(7.8)	Radiation Belt Interference	29	0.4	(6.2)
Morning Ramps	28	0.4	(4.0)	Poleward Boundary	18	0.3	(3.8)
Mincr Problems 1. Ambiguous (34) 2. Algorithm Failure (13) 3. Poleward Boundary (11) 4. Hand Error (1)	59	6.0	(8.5)	Minor Problems: 1: Ambiguous (12) 2. Algorithm Failure (2) 3. Hand Error (5)	10	0.3	(4.0)
Total Number of Morning Dis	Discrepancies > ± 3°:	> # 3°:	694	Total Number of Evening Discrepancies > ±	Discrepanc	ies > ± 2°:	47.1
% of Total Morning Boundaries with > 3° CGM Discrepancy:	es		10.8%	% of Total Evening Boundaries with > 2° CGM Discrepancy:	aries cy:		6.7%
(6398 Morning Boundaries Me	Measured)			(7030 Evening Boundaries Measured)	Measured)		

Evening ramps refer to the gradual onsets of precipitation in the evening sector. They occur in less than half of the passes and usually extend for less than half a degree in CGL. The onset of the evening ramps is chosen as the "better" boundary for the hand-chosen data set, but the algorithm tends not to pick up the onset of these evening ramps making the algorithm boundary closer to the 10^7 level. Referring to Table 4, it is seen that evening ramps account for nearly 2/3 of the $>\pm$ 2° CGL discrepancies. Actually, because in these cases the hand boundary is lower than the computer boundary they fall only in the negative direction. Since evening ramps are generally smaller than 2°, it is possible to deduce that the small regions of gradual onset of precipitation on the evening side are a major cause of the slightly lower hand-chosen evening boundary, and the negative skew seen in Figure 9.

Large irregularities on or near the equatorward edge of both the morning or evening auroral oval occur primarily under three conditions: (a) the occurrence of polar cap absorption events, (b) very low activity, and (c) abrupt changes in the activity. Polar cap absorption events have been, as much as possible, edited out of both data sets because they totally obscure the onset of auroral precipitation. Figure 5 illustrates ambiguous conditions on the evening-side equatorward edge of the aurora. The dashed line is the hand-chosen boundary, the dot-dash line the computer-chosen boundary. As was mentioned in Section 3, spurs are excluded from the hand-chosen data set because they may be produced by the precipitation of trapped electrons from prior activity. The algorithm usually picks the auroral boundary below the spur if it is broad enough, as was the case in this example. The net effect on the evening side of the > ± 2° irregularities is not significant in the distribution of differences shown in Figure 9, as 30 of the algorithm boundaries were lower and 33 of the hand boundaries were lower. However, this category of discrepancies can account for some of the very large differences between the two data sets.

Irregularities in the equatorward edge of the aurora produce a quite different net effect on the morning side. First, they account for more than half the $>\pm$ 3° CGL discrepancies on the morning side. Second, there is a distinct bias in that for more than 210 of these cases, the computer boundary is lower then the hand boundary. This category, then, accounts for many of the >+ 3° CGL differences in the distribution shown in Figure 10.

Radiation belt particles interfere with both morning and evening auroral boundaries. This accounts for ~30 percent of the large discrepancies on the morning side. (On the evening side the problem is minor, 0.4 percent of the large discrepancies, either because the sharp increase due to auroral precipitation is visible through the radiation belt signature or the generally higher latitude of the onset of auroral precipitation on the evening side produces a boundary above the latitude of particle radiation.) Radiation belt particle interference is a major problem in determining

equatorward auroral boundaries both by hand and with the computer and was discussed thoroughly in the previous sections. Although every attempt was made in both data sets to exclude radiation belt particles effects, it was not always successful. The effect is that the hand chosen boundaries in 200 of 270 cases where radiation particle effects are important, are lower than those chosen by the algorithm. On the morning side then, this category of discrepancy accounts for many of the > - 3° CGL differences in the distribution shown in Figure 10.

The category labelled Spurs, in Table 4, results from the triggering of the algorithm test too early by regions of enhancement which in these cases are clearly outside the auroral oval. This category accounts for many of the largest differences between the algorithm- and hand-chosen boundaries in the positive direction in Figures 9, 10 and 11. This is an algorithm failure which occurs for about 0.8 percent of the boundaries in the total data set.

The area of the aurora around noon local time is confusing, as was discussed earlier. Whereas, the method used to choose boundaries in this region by hand is fairly clearly defined and consistent, algorithm tests often fail to pick up the irregular conditions of the diffuse aurora. Instead the "cusp" boundary is chosen. An attempt has been made to edit out the algorithm "cusp" boundaries. The cases appear in the listing as the hand-chosen boundaries with no corresponding algorithm boundary. There are also instances in the noon sector where the algorithm chose boundaries which were not evident in the particle data. Because DMSP/F2 is in a dusk-dawn orbit this is not a major problem. If a satellite in a noon-midnight orbit is used to produce future Auroral Boundary Index values, further work on both the algorithm and hand selection of boundaries in the noon sector will be needed.

The remaining categories of large discrepancies between the two data sets are minor problems. The total number of boundaries in these additional cases accounts for less than 1 percent of the total boundary measurements. Morning ramps, which are a major problem in determining boundaries by hand, cause very few large discrepancies between the two data sets unlike their evening counterparts. Briefly, the other minor problems are: Poleward boundaries, the result of the algorithm tests being triggered very late (at or near the poleward edge of the aurora) because the gradient is very smooth; Ambiguous, where the reason for the discrepancy is a combination of two or more categories; Algorithm Failure, where the algorithm is wrong, but the reason is unclear; and Hand Errors, two of which were ephemeris interpolation errors, and the rest recording mistakes.

6. AURORAL BOUNDARY INDEX

A proposed Auroral Boundary Index was constructed in the following manner. First, using values of α and β from Table 1 a predicted Kp, called Kp¹, was calculated for each measured equatorward auroral boundary, λ_E , based on the relationship of λ and Kp shown in Eq. (1), using the following expression:

$$Kp^{\dagger} \equiv \frac{\lambda_{E} - \alpha_{i}}{\beta_{i}} . \tag{2}$$

Here α_i and β_i are the intercept and slope for the MLT zone in which the boundary was measured.

Second, using the Kp' calculated in Eq. (2), λ_{EM} , the projection of λ_{E} to the midnight (23-24) MLT sector, was calculated using the expression:

$$\lambda_{\rm EM}^{\ \prime} = \alpha_{\rm M}^{\ } + \beta_{\rm M}^{\ } \, \rm Kp^{\prime} \, , \tag{3}$$

where $\alpha_{\overline{M}}$ and $\beta_{\overline{M}}$ are the intercept and slope for 23-24 MLT zone. The last two columns in computer listing in Appendix B give the values of $\lambda_{\overline{EM}}$ ' for the handand computer-chosen values of $\lambda_{\overline{E}}$, respectively. Software has been developed to generate monthly plots of Kp', 3 hour averages of Kp' ($\overline{\text{Kp'}}$), and $\lambda_{\overline{EM}}$ ' for both the handand algorithm-chosen boundaries. Additionally, the derived quantities are plotted using all points and using morning or evening boundaries separately. A sample set of plots of $\overline{\text{Kp'}}$ and $\lambda_{\overline{EM}}$ ' for March 1978 is shown in Figures 12a-121.

As the final part of the development of the index we have correlated Kp with \overline{Kp}' for both hand- and algorithm-boundaries. The correlation was done using values of \overline{Kp}' derived from the morning sector boundaries, evening sector boundaries, and all boundaries. Table 5 shows the correlations of Kp vs \overline{Kp}' for each month in 1978. It is clear from Table 5 that correlations are best using only the evening boundaries (average of 0.84, hand and algorithm) and still quite good for the combined morning and evening boundary set (average of 0.84, hand, and 0.82, algorithm). The poorest correlations result when only morning boundaries are used (average of 0.77, hand, and 0.72, algorithm), a reflection of the morning problems. The lowest correlations generally occur in those months where there are very few active periods resulting in a restricted range for the data. Scatter plots of Kp vs of \overline{Kp}' for March 1978, are given in Figures 13a-13f. Straight lines show the results of performing linear regressions on Kp vs \overline{Kp}' and \overline{Kp}' vs Kp.

In order to assess the consistency between the hand-chosen and algorithm-chosen data sets, monthly regressions of $\overline{Kp'}$ derived from the hand-chosen boundaries on $\overline{Kp'}$ derived from the algorithm-chosen boundaries were calculated. The results of these regressions for March 1978 are shown in Figure 14. Average correlation coefficients of $\overline{Kp'}$ (hand) vs $\overline{Kp'}$ (algorithm) are 0.97 for evening values, 0.91 for morning values and 0.96 for combined values.

Table 5. Correlations of Kp vs Kp'

F==						
1973	Algorithm Kp' M/E	Algorithm Kp¹ Evening	Algorithm Kp¹ Morning	Hand Kp' M/E	Hand Kp¹ Evening	Hand Kp' Morning
Jan	0.87	0.86	0.80	0.89	0.86	0.85
Feb	0.74	0.78	0.64	0.77	0.77	0.66
Mar	0.81	0.84	0.73	0.83	0.84	0.75
Apr	0.81	0.83	0.70	0.84	0.84	0.76
May	0.85	0.37	0.77	0.88	0.87	0.84
Jun	0.78	0.81	0.67	0.77	0.80	0.66
Jul	0.76	0.78	0.68	0.79	0.79	0.73
Aug	0.86	0.87	0.83	0.89	0.89	0.83
Sep	0.87	0.89	0.76	0.89	0.90	0.83
Oct	0.78	0.81	0.66	0.80	0.82	0.73
Nov	0.85	0.86	0.74	0.87	0.87	0.81
Dec	0.83	0.84	0.73	0.86	0.86	0.79

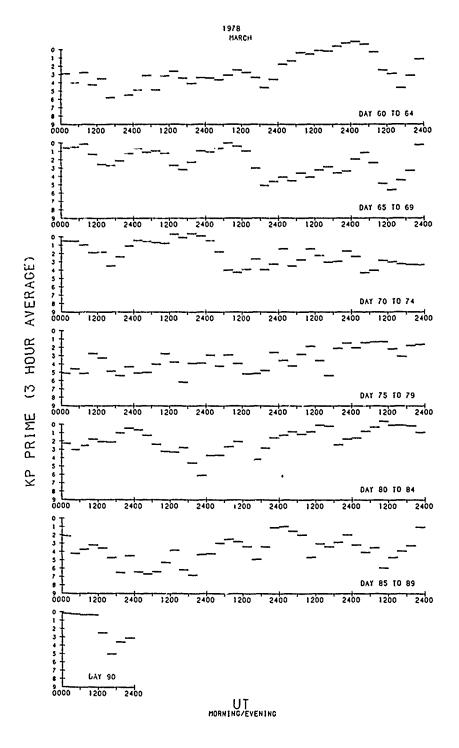


Figure 12a. Three-hour Average of Kp' Calculated From Hand-chosen Morning and Evening Equatorward Auroral Boundaries, Plotted as a Function of Universal Time for March 1978

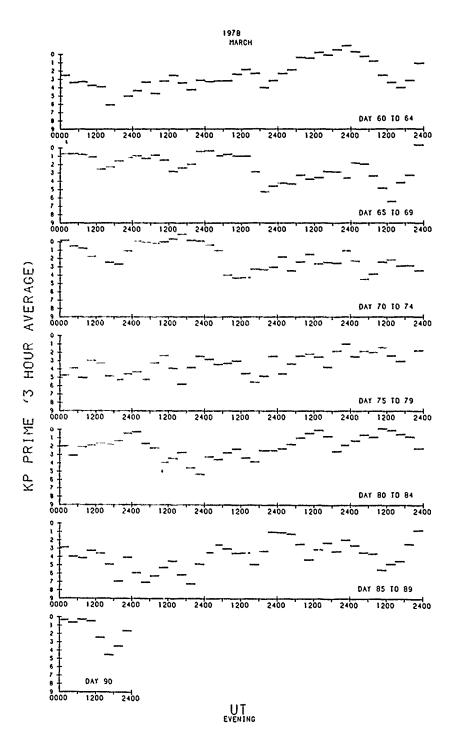


Figure 12b. Same as Figure 12a, Using Only Evening Boundaries

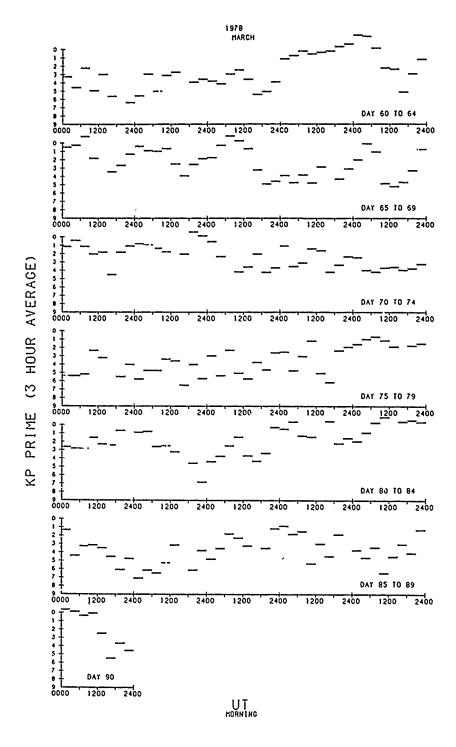


Figure 12c. Same as Figure 12a, Using Only Morning Boundaries

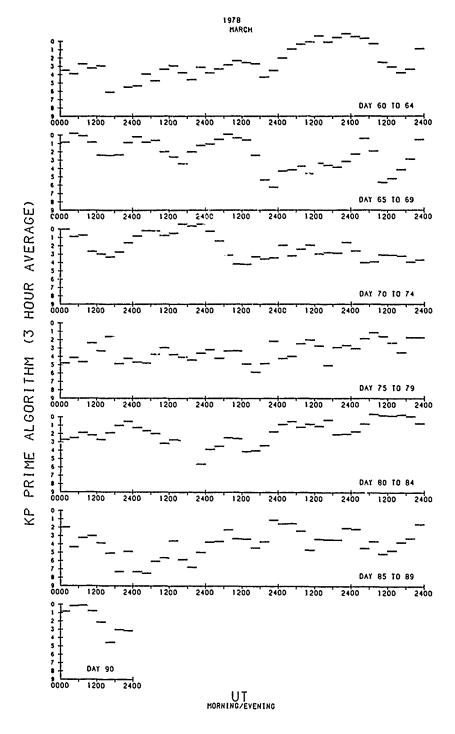


Figure 12d. Three-hour Average of Kp^1 Calculated From Algorithm-chosen Morning and Evening Equatorward Auroral Boundaries Plotted as a Function of Universal Time for March 1978

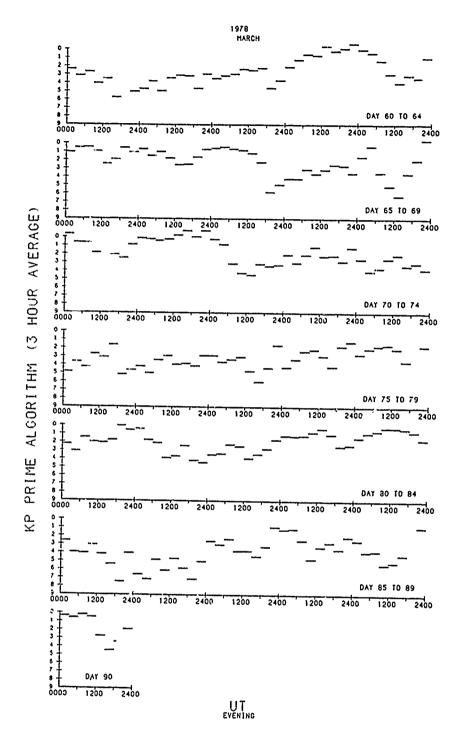


Figure 12e. Same as Figure 12d, Using Only Evening Boundaries

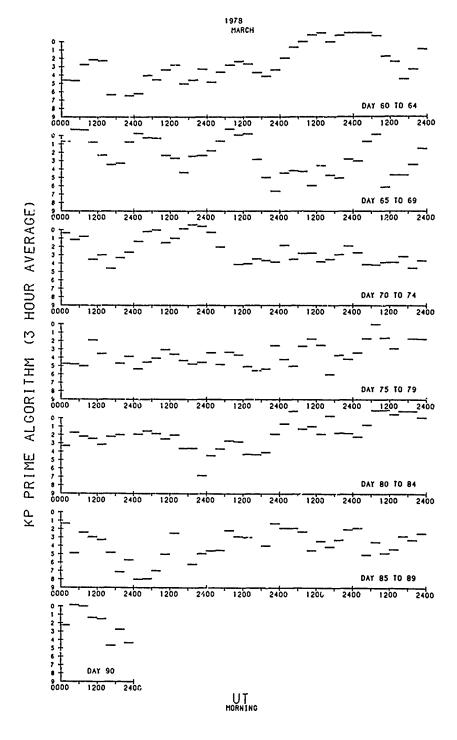


Figure 12f. Same as Figure 12d, Using Only Morning Boundaries

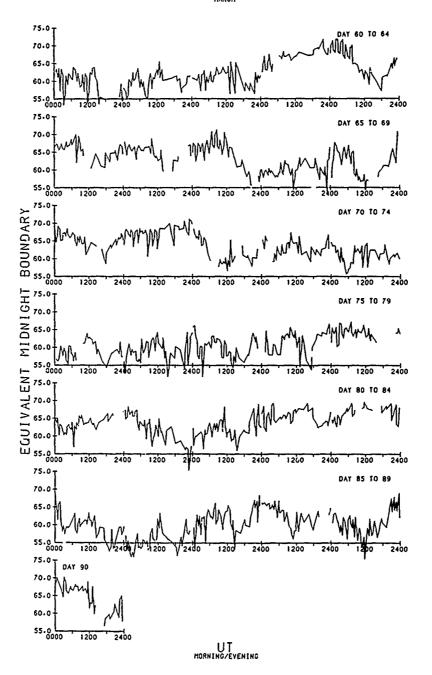


Figure 12g. Equivalent Midnight Boundary Calculated From Hand-chosen Morning and Evening Equatorward Auroral Boundaries, Plotted as a Function of Universal Time for March 1978

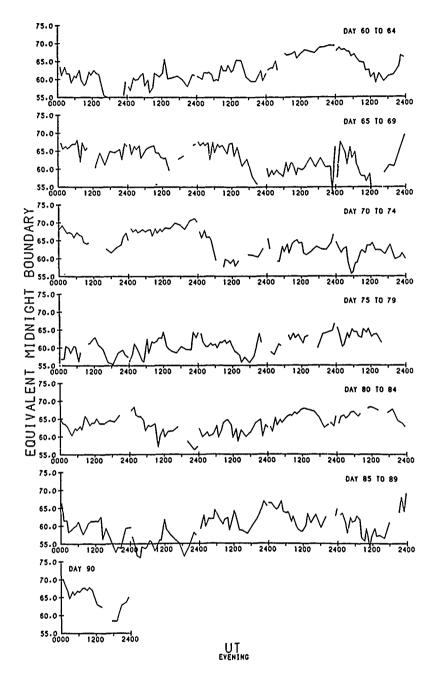


Figure 12h. Same as Figure 12g, Using Only Evening Boundaries

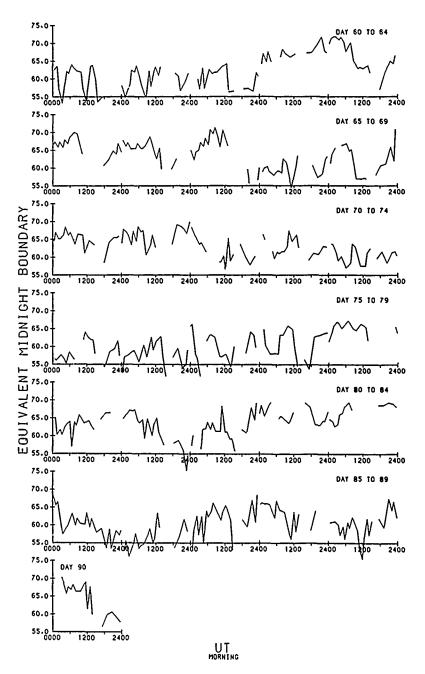


Figure 12i. Same as Figure 12g, Using Only Morning Boundaries

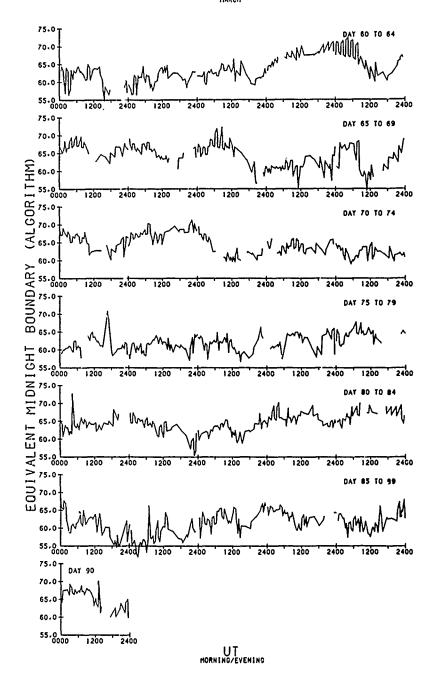


Figure 12j. Equivalent Midnight Boundary Calculated From Algorithmchosen Morning and Evening Equatorward Auroral Boundaries, Plotted as a Function of Universal Time for March 1978

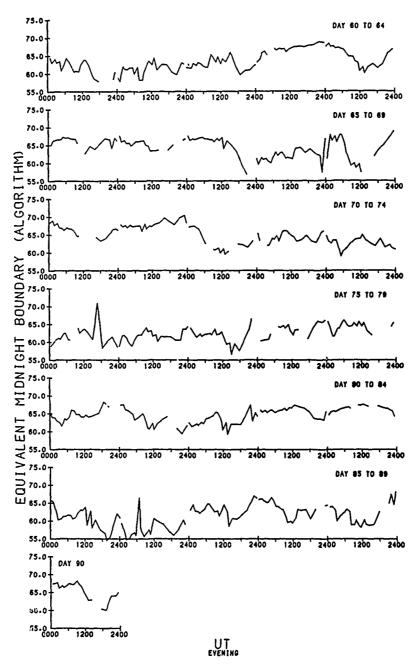


Figure 12k. Same as Figure 12j, Using Only Evening Boundaries



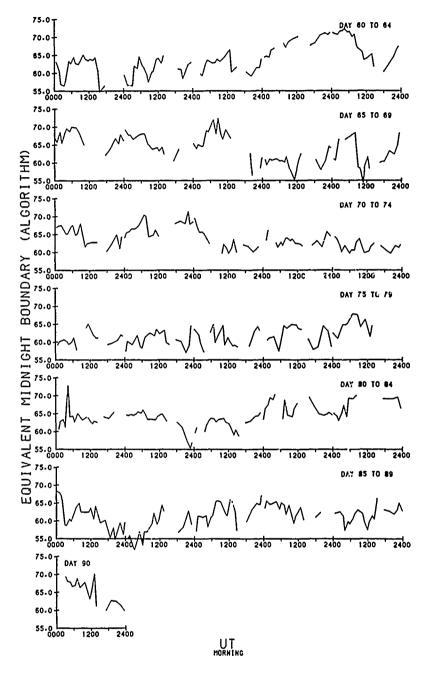


Figure 121. Same as Figure 12j, Using Only Morning Boundaries

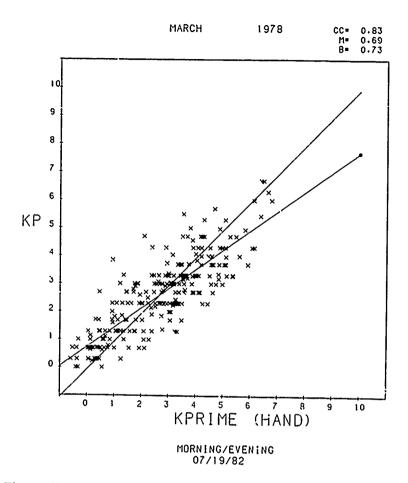


Figure 13a. Scatter Plot of the Values of Kp vs the Corresponding Three-hour Average Value of Kp' (labelled KPRIME) When Kp' is Calculated Using Hand-chosen Morning and Evening Equatorward Auroral Boundaries Obtained in March 1978. The result of performing a linear regression of Kp (Kp') vs Kp' (Kp) is shown by the straight line ending in a dot (not ending in a dot). The correlation coefficient, slope and intercept of the Kp vs Kp' line are shown in the upper right hand corner

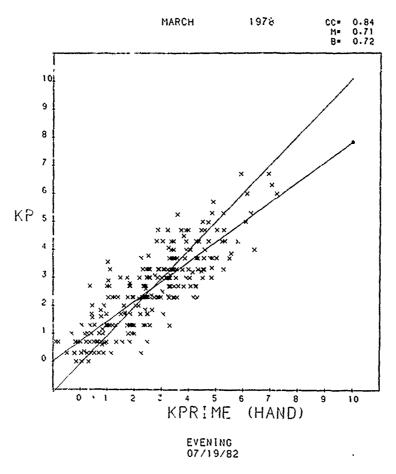


Figure 13b. Same as Figure 13a, Using Only Values of Kp¹ Determined From Evening Boundaries

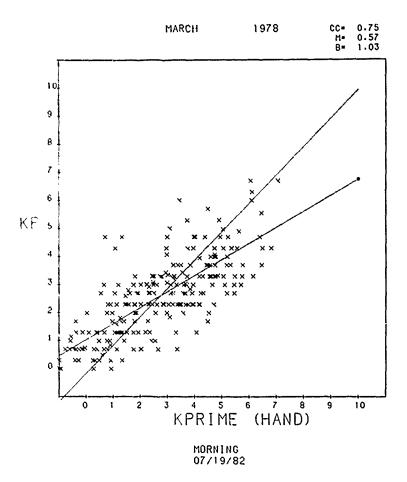


Figure 13c. Same as Figure 13a, Using Only Values of Kp¹ Determined From Morning Boundaries

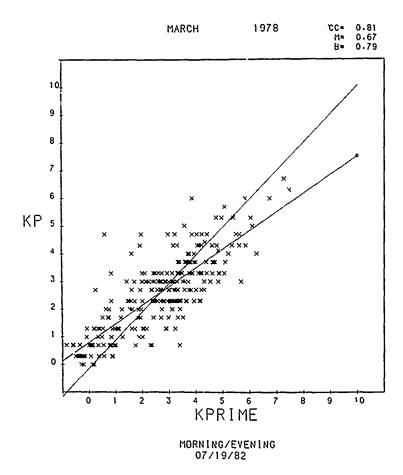


Figure 13d. Scatter Plot of the Values of Kp vs the Corresponding Three-hour Average Value of Kp' (labelled KPRIME) When Kp' is Calculated Using Computer-chosen Morning and Evening Equatorward Auroral Boundaries Obtained in March 1978. The result of performing a linear regression of Kp (Kp') vs Kp' (Kp) is shown by a straight line ending in a dot (not ending in a dot). The correlation coefficient, slope and intercept of the Kp vs Kp' line are shown in the upper right hand corner

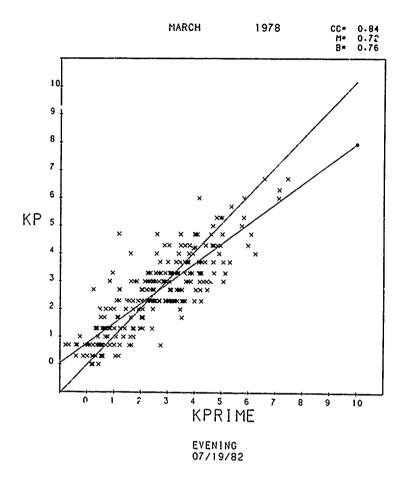


Figure 13e. Same as Figure 13d, Using Only Values of Kp' Determined From Morning Boundaries

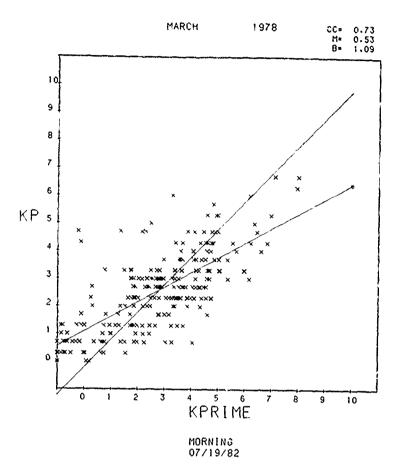


Figure 13f. Same as Figure 13d, Using Only Values of Kp' Determined From Morning Boundaries

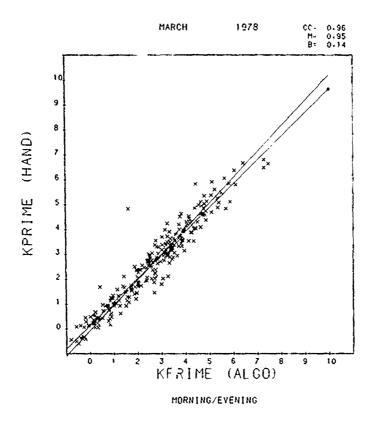


Figure 14. Scatter Plot of the Corresponding Values of Three-hour Averages of Kp' Calculated Using Hand-chosen [labelled KPRIME (HAND)] and Algorithm-chosen [labelled KPRIME (ALGO)] Morning and Evening Equatorward Boundaries Obtained in March 1978. The result of performing a linear regression of Kp' (hand) [Kp' (algorithm)] vs Kp' (algorithm) [Kp' (hand)] is shown by the straight line ending in a dot [not ending in a dot]. The correlation coefficient, slope and intercept of the Kp' (hand) vs Kp' (algorithm) line are shown in the upper right hand corner

7. DISCUSSION

Before making a comparison of the two sets of 1978 auroral equatorward boundaries, those laboriously done by hand and those done by a rather complex computer algorithm, we list the following overall conclusions:

(a) The boundaries cann be judged absolutely. There is a margin of ambiguity in all but the most abrupt and clean evening side rises of precipitating electron counts above background. In most cases the margin of ambiguity is small.

- (b) The evening boundaries have a much smaller margin of ambiguity than the morning boundaries. Characteristically the morning electron precipitation onset is different from that of the evening, having long ramps and more irregularities. This may well reflect a difference in the dynamics of the plasma sheet at different local times and/or a difference in precipitation mechanisms. The evening boundaries also have higher correlation coefficients than the morning boundaries, when directly related to Kp.
- (c) Contamination of the auroral electron signal by radiation belt particles presents the single greatest problem in choosing equatorward boundaries by both methods. Again, this is particularly true on the morning side. To avoid this problem on future DMSP flights additional shielding has been added to the detectors.

Since one of the objects of this study is an assessment of the algorithm developed by Hardy and Holeman⁸ for choosing boundaries we make the following additional conclusions from comparing the two boundary sets:

- (d) Both hand- and computer-chosen boundary sets have approximately the same internal consistency. The spread in boundary values from the linear regression values are very similar (Figures 7 and 8).
- (e) Although the two sets of boundaries result in different linear regressions (systematically differing slopes and intercepts) they correlate equally well with Kp (Table 1).
- (f) The two boundary sets show similar discrepancies between morning and evening boundaries, the latter being more consistent both within and between boundary sets and having high correlation with Kp.
- (g) When used as a predictor of Kp (using the three-hour averaged values of Kp¹) the accuracy of the two sets is indistinguishable when the evening boundaries are used, and is only slightly better for hand-chosen boundaries using morning and combined values (Table 5). This is clearly shown in Figure 14 in which the two predicted values agree with a correlation coefficient of 0.96.

We therefore conclude that the auroral equatorward boundaries selected by the algorithm described here may be used for both scientific and predictive purposes in place of the hand-chosen boundaries. The algorithm was developed and tested for dawn and dusk boundaries and is not applicable in the noon sector.

Finally, we use the auroral equatorward boundaries as a measure of auroral activity by indexing the activity with the boundary measurements. Using the algorithm for choosing boundaries at the site of the down-link data transmission without intermediate processing can give the index in near real time. The most direct index is simply the boundary itself, tagged by both universal and local time. The large offset of the auroral oval in magnetic coordinates toward higher latitudes at noon and lower latitudes at midnight, makes the use of multi-local time values of

the boundary difficult at best, and no doubt misleading to casual users. We therefore choose to scale each boundary to an equivalent midnight boundary by way of the derived Kp = Kp'. Kp' is not used as the index for two reasons: (a) the index is derived from boundaries, not from magnetic activity. The source of the index is obscured by using Kp'. (b) The boundary measurements have a finer time resolution, and one more appropriate to substorm activity than the 3-hr Kp interval. A finer Kp' resolution would again add confusio. in interpretation.

Another choice remains. The equivalent midnight, $\lambda_{\rm EM}$ used for the index can be derived from both morning and evening boundaries, or either morning or evening boundaries. Plots of the combined data (Figure 12g) have better time resolution, but contain a great deal of variation from point to point (oscillating from a morning to an evening equivalent boundary). These variations result either from the poorer quality of morning boundary choices and/or from a real difference in the particle dynamics between the two local time sectors. Because the evening boundary set has fewer and more controllable selection problems, is more internally consistent, and correlates more directly with Kp, we choose the equivalent midnight boundary obtained from the evening sector boundary set (Figure 12h) as the Auroral Boundary Index. To reiterate, the Auroral Boundary Index is the projected midnight equatorward auroral boundary found using an actual evening sector boundary and the statistically determined systematic local time variation of the oval. During periods of good data accumulation there will be one index value each 55 min. The Auroral Boundary Index is presented month by month for 1978 in Appendix A.

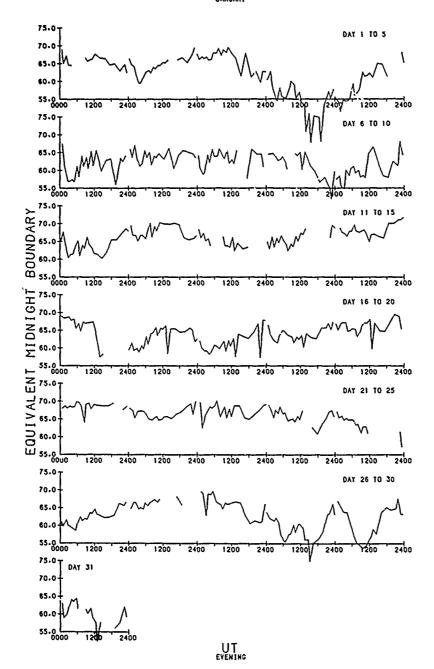
References

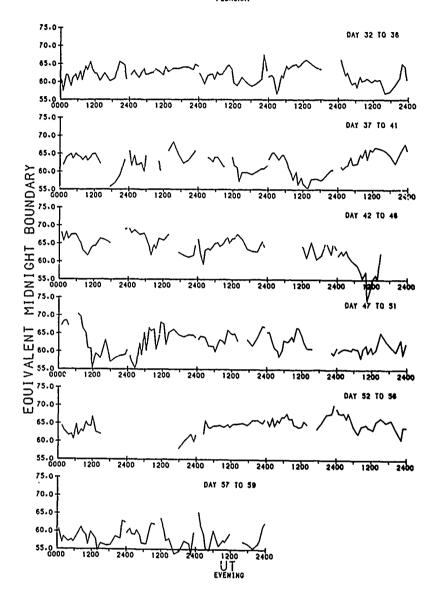
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- 5. Slater, D.W., Smith, L.L., and Kleckner, E.W. (1980) Correlated observations of the equatorward diffuse auroral boundary, J. Geophys. Res. 85:531.
- 6. Gussenhoven, M.D., Hardy, D.A., and Burke, W.J. (1981) DMSP/F2 electron observations of equatorward auroral boundaries and their relationship to magnetospheric electric fields, J. Geophys. Res. 86:768.
- 7. Hardy, D.A., Burke, W.J., Gussenhoven, M.S., Heinemann, N., and Holeman, E. (1981) DMSP/F2 electron observations of equatorward auroral boundaries and their relationship to the solar wind velocity and the north-south component of the interplanetary magnetic field, J. Geophys. Res. 86:9961.
- 8. Hardy, D. A., and Holeman, E. (1983) The Global Auroral Boundary Code for the Global Weather Central of the Air Weather Service (to be published).
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- 10. Hardy, D.A., Gussenhoven, M.S., and Huber, A. (1979) The Precipitating Electron Detectors (SSJ/3) for the Block 5D/Flights 2-5 DMSP Satellites: Calibration and Data Presentation, AFGL-TR-79-0210, AD A083136.

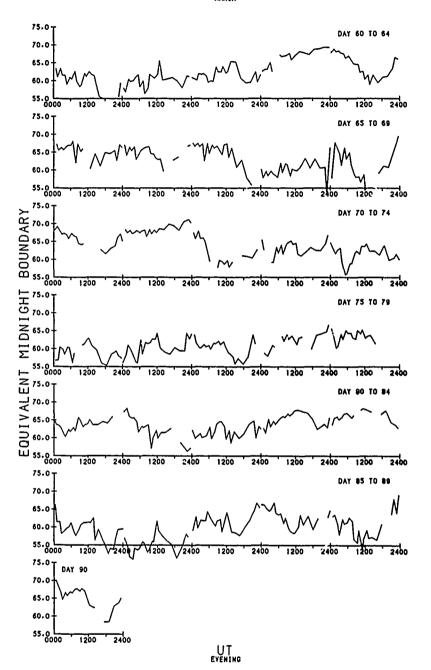
Appendix A

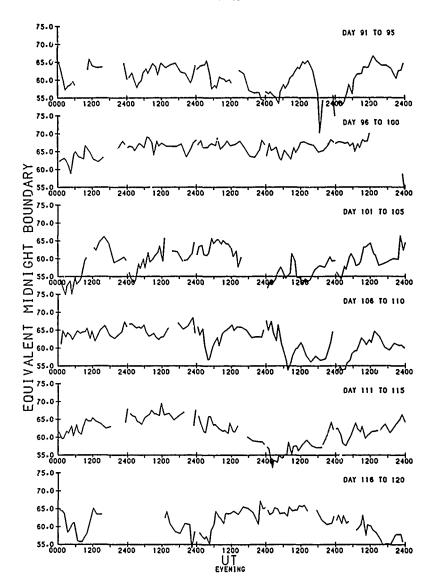
Auroral Boundary Index

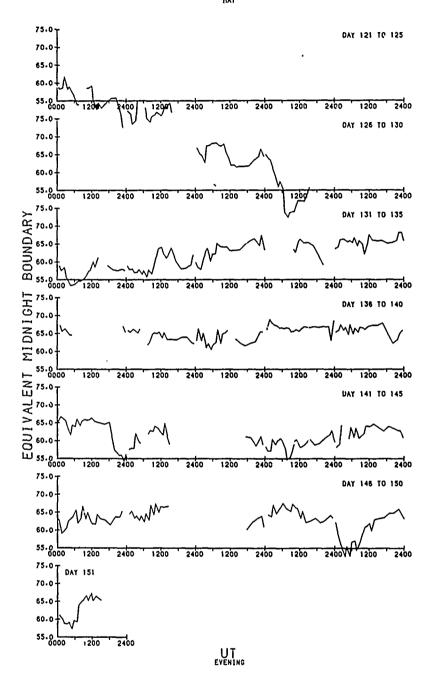
Each value of the Auroral Boundary Index is a projected midnight equatorward auroral boundary found using an actual evening sector boundary and the statistically determined systematic local time variation of the auroral oval. The Index is presented in the following pages as a function of universal time, by month, for 1978.

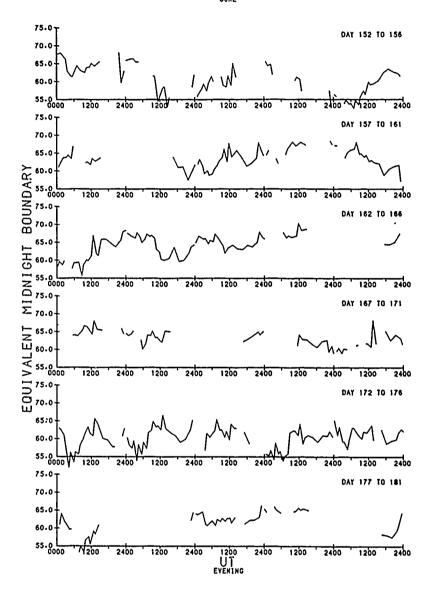


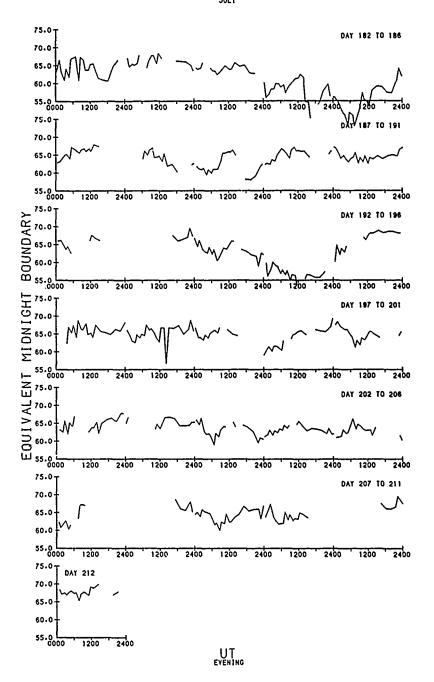


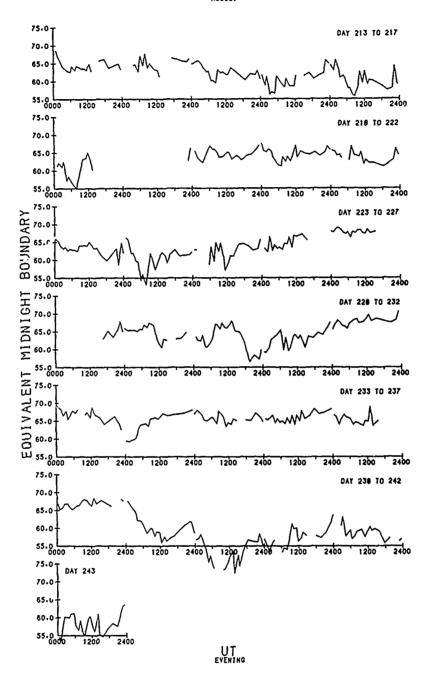




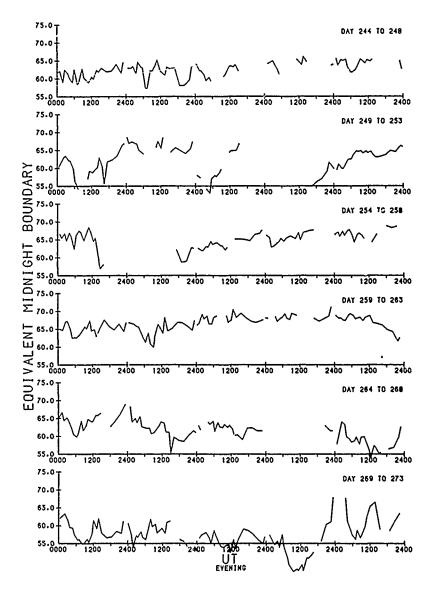


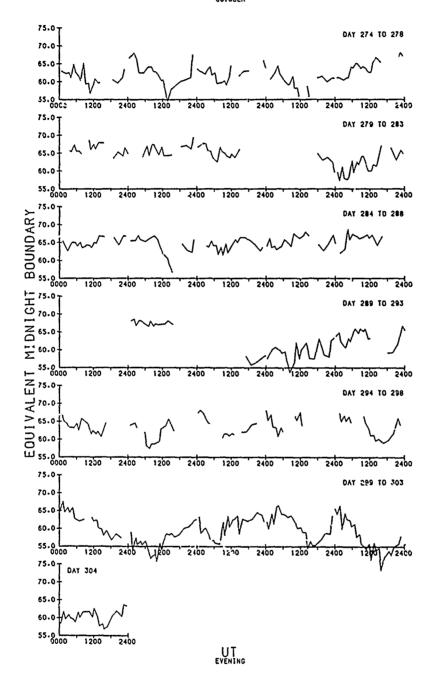




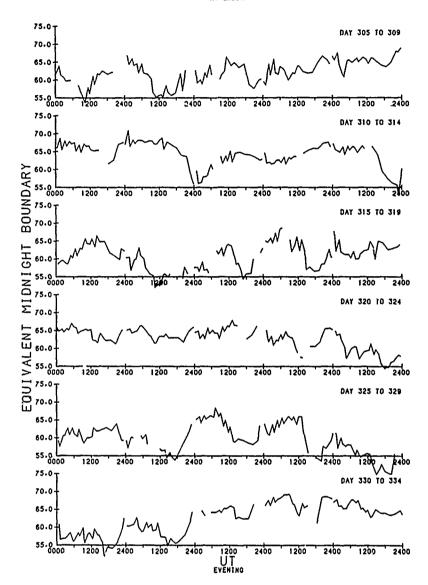




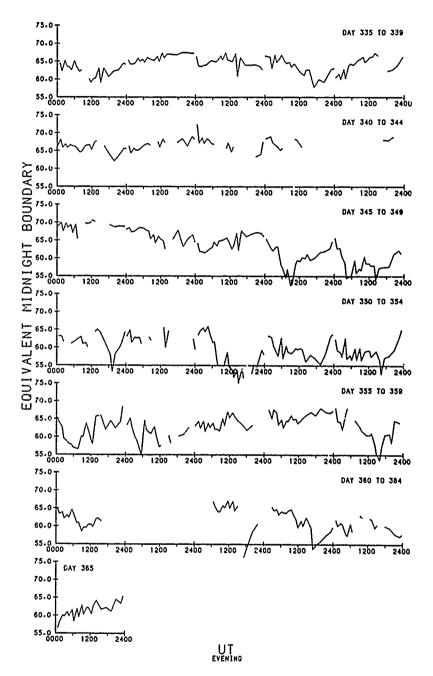




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Appendix B

1978 Auroral Boundary Listing

The following computer listing contains equatorward auroral boundary measurements made by hand and computer from DMSP/F2 precipitating electron particle data for the entire year of 1978. There are two sets of vertical columns on a page. A single set reads from left to right.

YR = 1978

DA = Julian Day

HDDK/HR, MN, SEC = Exact universal time at which the hand-chosen

boundary was measured.

PO = Pole

HDDK/LT* = Exact magnetic local time in hours at which the

hand-chosen boundary was measured.

HDDK/LAT* = Exact corrected geomagnetic latitude at which.

the hand-chosen boundary was measured.

CDDK/LT* = Exact magnetic local time in hours at which

the computer-chosen boundary was measured.

CDDK/LAT* = Exact corrected geomagnetic latitude at which

the computer-chosen boundary was measured.

DLAT* = Computer-chosen boundary minus hand-chosen boundary.

MIDNIGHT/LATH*

= Projection of the hand-chosen boundary (HDDK/LAT) measured in the given MLT (HDDK/LT) to the corrected geomagnetic latitude at which the boundary would occur in the midnight (23-24) magnetic local time bin, calculated as described in Section 5 of the preceding paper.

MIDNIGHT/LATC*

= Same as MIDNIGHT/LATH except using CDDK/LAT and CDDK/LT.

*To read values in these columns correctly, multiply by 10⁻¹.

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